

How accurate is it to update the cosmology of your halo catalogues?

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ABSTRACT

We test and present the application of the full rescaling method by Angulo & White (2010) to change the cosmology of halo catalogues in numerical simulations for cosmological parameter search using semi-analytic galaxy properties. We show that a reduced form of the method can be applied in small simulations with box side of $\sim 50 \text{ h}^{-1}\text{Mpc}$ or smaller without loss of accuracy. We perform statistical tests on the accuracy of the properties of rescaled individual haloes, and also on the rescaled population as a whole. We find that individual positions and velocities are recovered with almost no detectable biases, but with a scatter that increases slightly with the size of the simulation box when using the full method. The dispersion in the recovered halo mass does not seem to depend on the resolution of the simulation. Regardless of the halo mass, the individual accretion histories, spin parameter evolution and fraction of mass in substructures are remarkably well recovered. In particular, in order to obtain a more accurate estimate of the halo virial mass, it was necessary to apply an additional correction due to the change of the virial overdensity and the estimate of its effect on a NFW virial mass. The mass of rescaled haloes can be underestimated (overestimated) for negative (positive) variations of either σ_8 or Ω_m , in a way that does not depend on the halo mass. Statistics of abundances and correlation functions of haloes show also small biases of < 10 percent when moving away from the base simulation by up to 2 times the uncertainty in the WMAP7 cosmological parameters. The merger tree properties related to the final galaxy population in haloes also show small biases; the time since the last major merger, the assembly time-scale, and a time-scale related to the stellar ages show correlated biases which indicate that the spectral shapes of galaxies would only be affected by global age changes of $\sim 150\text{Myr}$, i.e. relatively small shifts in their broad-band colours. We show some of these biases for different separations in the cosmological parameters with respect to the desired cosmology so that these can be used to estimate the expected accuracy of the resulting halo population. We also present a way to construct grids of simulations to provide stable accuracy across the Ω_m vs. σ_8 parameter space.

Key words: cosmology: cosmological parameters - cosmology: large-scale structure - cosmology: theory - methods: numerical

1 INTRODUCTION

The Lambda Cold Dark Matter (Λ CDM) cosmological model is the standard theoretical framework for structure

formation in the Universe. In order to understand how galaxies form and evolve in this cosmological context, we must also understand the properties of dark matter haloes over a wide range of physical scales and across the cosmic history. Numerical simulations provide one of the best methods for approaching this problem.

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The cosmological parameters that provide the best match between the Λ CDM cosmology and several observations, are most often obtained using measurements of the power spectrum of temperature fluctuations in the cosmic microwave background (Sánchez et al. 2006; Hinshaw et al. 2009; Dunkley et al. 2009; Jarosik et al. 2010) or of density fluctuations in the galaxy distribution (Sánchez et al. 2006; Percival et al. 2007); but these are limited to the linear regime of density fluctuations. In order to make comparisons between the model and observations in the non-linear regime, without making simplified assumptions such as the spherical or ellipsoidal collapse (Sheth, Mo & Tormen 2001), one would need to use fully non-linear numerical simulations. However, these are too expensive in terms of computational time. Great efforts go into running even single simulations corresponding to one set of cosmological parameters, which can still be compared with a very wide range of observational measurements.

For the past several years the Millennium Simulation (Springel et al. 2005) has been the focus for many studies of the distribution and statistical properties of dark matter haloes and provides the basis for the implementation of semi-analytic models of the evolving galaxy population (Bower et al. 2006; Croton et al. 2006). However, it is important to appreciate that this simulation was run using the WMAP1 cosmological parameter set (Spergel et al. 2003), which are rather different from the current best fit parameters. More recently the Bolshoi Simulation (Klypin, Trujillo-Gomez & Primack 2010) used cosmological parameters consistent with the latest WMAP5 (Hinshaw et al. 2009; Komatsu et al. 2009; Dunkley et al. 2009) and WMAP7 (Jarosik et al. 2010) values. The main difference is that the Millennium simulation used a substantially larger amplitude of perturbations than the Bolshoi one.

The differences in the parameters of the cosmological model can affect the abundances and properties of the dark matter haloes at a given redshift. Consequently, the properties of the galaxies hosted by such haloes may also be affected. As a consequence the interpretation of many observational statistical tests, like those measuring clustering and luminosity functions, becomes more difficult. Notice that according to the cosmological parameter constraints from WMAP5 and cluster abundances (Rozo et al. 2009), the Millennium Simulation is about $2 - \sigma$ away from the best fit model (see Figure 1 in Klypin, Trujillo-Gomez & Primack 2010). The small variations of the parameters of the background cosmological model also become important in the understanding of the baryonic processes involved in galaxy formation. An extreme case can be seen in Cole et al. (1994) where the maximum of the stellar formation rate (SFR) activity is found at $z = 1$ due to their choice of $\Omega = 1$, which favours mergers at low redshifts.

Many of the differences between the Millennium and Bolshoi simulations can be re-scaled using different methods. A first attempt to change the cosmology of a simulation was presented by Zheng et al. (2002) which was later applied by Harker, Cole & Jenkins (2007) with the aim of constraining the cosmological parameter related to the amplitude of linear fluctuations in spheres of $8h^{-1}\text{Mpc}$, σ_8 , using the GALFORM semi-analytic model (Baugh et al. 2005) and only two individual dark matter simulations. This approach already

greatly diminished the computational time that would have involved running numerical simulations for each set of cosmological parameters. More recently, the method suggested by Angulo & White (2010, AW10) allows a more flexible change in the cosmological parameters of a simulation without incurring in an important loss of precision. This algorithm scales the output of a cosmological N -body simulation carried out for one specific set of cosmological parameters so that it represents the growth of structure in a different cosmology. The accuracy of the rescalings can only be estimated at first order using extended Press-Schechter theory or, more accurately, by running numerical simulations and testing for particular characteristics.

However, the algorithm developed by AW10 is applied to every particle in the simulation and the post-processing of the outputs of the simulation must be repeated (i.e. identification of the haloes and construction of the merger trees). This can be computationally demanding, particularly in the case when the rescaling needs to be done several times. We are interested in studying whether applying the method to dark-matter haloes instead of to particles produces similar accuracies in the rescaled simulations, since this would already considerably reduce the computational time.

We study whether the use of the AW10 method allows to explore variations in the phenomenology of non-linear density fluctuations in the cosmological parameter space via, for example, monte-carlo sampling. This is the reason why we will concentrate on reducing as much as possible the computational time in obtaining a catalogue of haloes from an N -body simulation on a different cosmology using AW10, and also on obtaining even a reduced version of their method. In any of these cases, the computational time involved is dramatically reduced in comparison to what would be needed to run a complete simulation for each new set of cosmological parameters. In the case of the reduced algorithm, it can be applied to small simulations where linear corrections span scales of the order of the box size. In the full and reduced cases, the approach we follow can be used when the particle information is not available as will be the case of future large simulations. We perform a number of tests using N -body simulations in similar cosmologies than those corresponding to the Millennium and Bolshoi simulations, in order to quantify the accuracy of the recovery of individual and statistical properties of the dark matter haloes and also exploring the effect of different extrapolation baselines.

The outline of this paper is the following. In Section 2 we present a brief description of the rescaling technique presented by AW10. We also describe in this section the N -body simulations used in this work, and the applicability of a reduced version of the method. The results of the statistical and individual properties of rescaled dark matter haloes are shown in Section 3. In Section 4 we measure how the resulting halo catalogues are affected by the distance in the cosmological parameter plane $\Omega_m - \sigma_8$ around the base WMAP7- Λ CDM model; this can later be used to evaluate how precise is any rescaled dark matter halo catalogue to ensure optimal explorations of the cosmological parameter space. Finally, the conclusions are presented in Section 5.

Table 1. Relevant parameters used in the simulations. The columns show the name of the simulation, the number of particles, the gravitational softening, the initial redshift, the particle mass, the boxsize, the matter density parameter, the baryon density parameter, the primordial spectral index, the dimensionless Hubble parameter, and the linear fluctuations amplitude in spheres of $8h^{-1}\text{Mpc}$. The cosmological constant density parameter is $\Omega_\Lambda = 1 - \Omega_m$ (flat models) in all cases. The boxsize of the B simulations were calculated using the resulting scaling factor s obtained from the rescaling technique.

Name	N_p	ϵ [$h^{-1}\text{kpc}$]	z_i	M_p [$h^{-1}\text{M}_\odot$]	L_{box} [$h^{-1}\text{Mpc}$]	Ω_m	Ω_b	n	h	σ_8
<i>A</i>	256^3	5	74.7	8.93×10^8	60.00	0.25	0.0450	1.00	0.73	0.90
<i>Alow</i>	128^3	15	59.7	7.15×10^9	60.00	0.25	0.0450	1.00	0.73	0.90
<i>Abig</i>	256^3	30	43.3	9.08×10^{10}	280.00	0.25	0.0450	1.00	0.73	0.90
<i>B</i>	256^3	5	61.2	1.38×10^9	67.68	0.27	0.0469	0.95	0.70	0.82
<i>Blow</i>	128^3	15	49.1	1.11×10^{10}	67.68	0.27	0.0469	0.95	0.70	0.82
<i>Bbig</i>	256^3	30	36.3	1.27×10^{11}	305.33	0.27	0.0469	0.95	0.70	0.82
<i>Bo+2</i> ($\Omega_m + 2\sigma$)	128^3	20	55.7	6.39×10^9	52.72	0.33	0.0469	0.95	0.70	0.82
<i>Bo+1</i> ($\Omega_m + 1\sigma$)	128^3	20	52.5	8.27×10^9	59.29	0.30	0.0469	0.95	0.70	0.82
<i>Bo-1</i> ($\Omega_m - 1\sigma$)	128^3	20	45.7	1.55×10^{10}	78.78	0.24	0.0469	0.95	0.70	0.82
<i>Bo-2</i> ($\Omega_m - 2\sigma$)	128^3	20	42.2	2.34×10^{10}	94.48	0.21	0.0469	0.95	0.70	0.82
<i>Bs+2</i> ($\sigma_8 + 2\sigma$)	128^3	20	52.8	1.11×10^{10}	67.68	0.27	0.0469	0.95	0.70	0.88
<i>Bs+1</i> ($\sigma_8 + 1\sigma$)	128^3	20	50.9	1.11×10^{10}	67.68	0.27	0.0469	0.95	0.70	0.85
<i>Bs-1</i> ($\sigma_8 - 1\sigma$)	128^3	20	47.3	1.11×10^{10}	67.68	0.27	0.0469	0.95	0.70	0.79
<i>Bs-2</i> ($\sigma_8 - 2\sigma$)	128^3	20	45.5	1.11×10^{10}	67.68	0.27	0.0469	0.95	0.70	0.76
<i>Bo+1s+1</i> ($\Omega_m + 1, \sigma_8 + 1\sigma$)	128^3	20	54.4	8.27×10^9	59.29	0.30	0.0469	0.95	0.70	0.85

2 SCALING THE HALO CATALOGUES OF N -BODY SIMULATIONS

We now briefly describe the procedure used to scale halo catalogues from a given cosmology into a new set of cosmological parameters. This procedure consists on the method presented by AW10 which can either be applied to haloes or individual particles in the simulation. If $P(k)$ is the linear matter power spectrum at $z = 0$ we can define the variance of the linear density field as,

$$\sigma^2(R, z) = \frac{D(z)^2}{4\pi} \int_0^\infty k^2 P(k) W^2(kR) dk, \quad (1)$$

where R is a comoving smoothing scale, $D(z)$ is the linear growth factor normalised so that $D(z = 0) = 1$ and $W(x)$ is the Fourier transform of a spherical top-hat filter defined by,

$$W(x) = 3 \frac{\sin(x) - x \cos(x)}{x^3}. \quad (2)$$

Assuming that we want a halo catalogue of a given cosmology (denoted with B) evolved to a final redshift z_B^f starting from another halo catalogue which has different cosmological parameters (denoted with A), the procedure is to find a length scaling s of the boxsize, and a final redshift z_A^f defined so that the linear fluctuation amplitude $\sigma_A(s^{-1}R, z_A^f)$ over the range $[s^{-1}R_1, s^{-1}R_2]$ is as close as possible to $\sigma_B(R, z_B^f)$ over the range $[R_1, R_2]$. This is done by minimising the function

$$\delta_{rms}^2 = \frac{1}{\ln(R_2/R_1)} \int_{R_1}^{R_2} \left[1 - \frac{\sigma_A(s^{-1}R, z_A^f)}{\sigma_B(R, z_B^f)} \right]^2 \frac{dR}{R}, \quad (3)$$

over s and z_A^f . The values of R_1 and R_2 are selected so that $M(s^{-1}R_2)$ is the mass of the largest halo in the original simulation and $M(s^{-1}R_1)$ the mass of the smallest one.

Once we have the two parameters s and z_A^f , the boxsize is scaled so that $L_B = sL_A$ and the earlier redshifts in the

new cosmology ($z_B < z_B^f$) are obtained from those in the original cosmology ($z_A < z_A^f$) through

$$D_B(z_B) = D_A(z_A) \frac{D_B(z_B^f)}{D_A(z_A^f)}. \quad (4)$$

After scaling the positions with s , and having matched the cosmological times, we must consider that the velocity and mass of the haloes need to be corrected using

$$\vec{v}_B = s \frac{(1+z_A)}{(1+z_B)} \frac{\dot{D}_B(z_B)}{\dot{D}_A(z_A)} \frac{h_A}{h_B} \vec{v}_A, \quad (5)$$

$$M_B = s^3 \frac{\Omega_B}{\Omega_A} M_A, \quad (6)$$

where the dot indicates a derivative with respect to time, h is the dimensionless Hubble parameter and Ω_B and Ω_A are the values of the matter density parameter $\Omega_m = \Omega_b + \Omega_{dm}$ for different cosmologies, being Ω_b and Ω_{dm} the baryon and dark matter density parameters, respectively.

The steps presented in equations (1) to (6) will be referred to as the reduced AW10 method. The full implementation requires to correct the positions and velocities for residual differences in the power spectrum of the two cosmologies and/or for the length scaling. To do this it is necessary to modify the contribution of the long wavelength components on the position and velocity fields using the Zel'dovich approximation. The range of modes where the correction is applied is $k < k_{nl}$, where k_{nl} satisfies $\Delta^2(k_{nl}) = 1$ being $\Delta^2(k) = k^3 P(k)/2\pi^2$ (See AW10 for the full details on this correction).

2.1 N -body simulations

In order to test the scaling of haloes and their histories we use two sets of simulations. In a first set we use 6 simulations. The two main simulations are simulation A which has

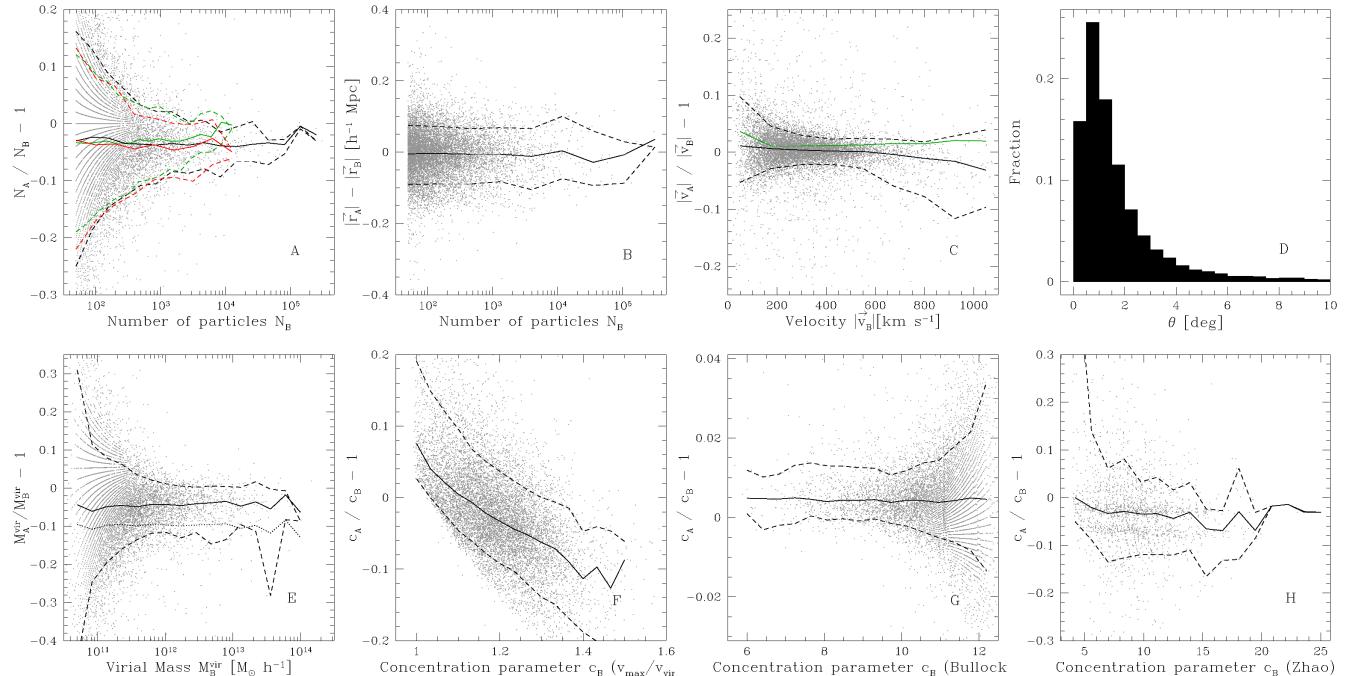


Figure 2. Comparison of individual properties between rescaled and reference haloes. The top boxes show the relative difference in number of particles (A), the difference between halo positions as a function of the number of particles per halo (B), the relative difference between the modulus of the velocity vectors as a function of the peculiar velocity of haloes (C) and the histogram of the angle subtended by the velocity vectors of matched haloes (D). The bottom boxes show the virial mass (E) and three different halo concentration parameters, obtained from the ratio v_{max}/v_{vir} (F), computed using the Bullock et al. (2001) recipe (G) and usig the Zhao et al. (2008) prescription (H). In all panels we show the full population of haloes. The black solid lines show the median, and black ashed lines enclose 80 percent of the haloes. We also show the recovered number of particles for the simulation with a large box and with a lower resolution (panel A, green and red lines respectively), the median velocity offset for the large simulation (panel C, green solid line) and the median of the relative error in the virial masses before the correction of $\delta_{vir}(z, \Omega_m)$ via a NFW profile (panel E, dotted line.)

Table 2. Rescaling parameters. In the first column we show the simulation name, the second and third contain the scaling factor s and the final redshift z_A^f , respectively, obtained after minimising Eq. 3. The last column shows the minimum δ_{rms}^2 resulting from minimising the difference between the rescaled and desired linear fluctuation amplitudes. With the exception of simulation *Bbig* where the range $[M_1, M_2] = [1 \times 10^{11}, 5 \times 10^{15}]h^{-1}M_\odot$ is used, in all simulations we use the mass range $[M_1, M_2] = [1 \times 10^{10}, 5 \times 10^{14}]h^{-1}M_\odot$ to perform the minimisation.

Name	s	z_A^f	δ_{rms}^2
<i>B</i>	1.128	0.361	2.6×10^{-5}
<i>Blow</i>	1.128	0.361	2.6×10^{-5}
<i>Bbig</i>	1.091	0.325	2.6×10^{-5}
<i>Bo</i> +2	0.779	-0.383	3.4×10^{-7}
<i>Bo</i> +1	0.876	-0.191	3.9×10^{-7}
<i>Bo</i> -1	1.164	0.201	9.5×10^{-8}
<i>Bo</i> -2	1.396	0.425	5.2×10^{-7}
<i>Bs</i> +2	1.000	-0.147	1.6×10^{-7}
<i>Bs</i> +1	1.000	-0.077	2.7×10^{-7}
<i>Bs</i> -1	1.000	0.076	5.3×10^{-8}
<i>Bs</i> -2	1.000	0.153	1.5×10^{-7}
<i>Bo</i> +1 <i>s</i> +1	0.876	-0.273	7.2×10^{-8}

a Millennium-like cosmology and simulation *B* which has the background cosmology of the Bolshoi simulation. Both simulations contain 256^3 particles in cubic volumes of 60

and $67.68h^{-1}\text{Mpc}$ of side length, respectively, which results in a particle mass comparable to that of the Millennium Simulation, $M_p \simeq 10^9h^{-1}M_\odot$. This set of simulations also includes a low resolution version of simulations *A* and *B*, with 128^3 particles, and versions with volumes ~ 100 times bigger (with respect to *A* and *B*), which we will refer to as *Alow*, *Blow* and *Abig*, *Bbig*, respectively. With the low resolution simulations we will minimise the rescaling function using the same parameters as for the *A* and *B* simulations; in this way we will only test the effect from the coarser resolution. In the case of the larger volume, the minimisation will be performed on a range of masses one order of magnitude higher than for *A* and *B* and will therefore include the effects from considering different scales in the power spectra, as well as a lower resolution.

Unless otherwise stated, the *B* cosmology will be taken as the reference model and the simulations with the *A* cosmology will be rescaled to resemble the former as closely as possible.

The second set contains 9 low resolution simulations (128^3 particles) designed to measure the variation in the precision of the rescaling of the halo catalogues as the desired cosmology moves further away in the plane $\Omega_m - \sigma_8$ from the original parameters. In order to do that, we fix all parameters and vary only Ω_m (or σ_8) in $\pm 1\sigma$ and $\pm 2\sigma$ of its original value using $\sigma \sim 0.03$, which is in concordance with the standard WMAP7 deviation for both parameters (Jarosik et al. 2010), giving us 8 simulations. The remaining

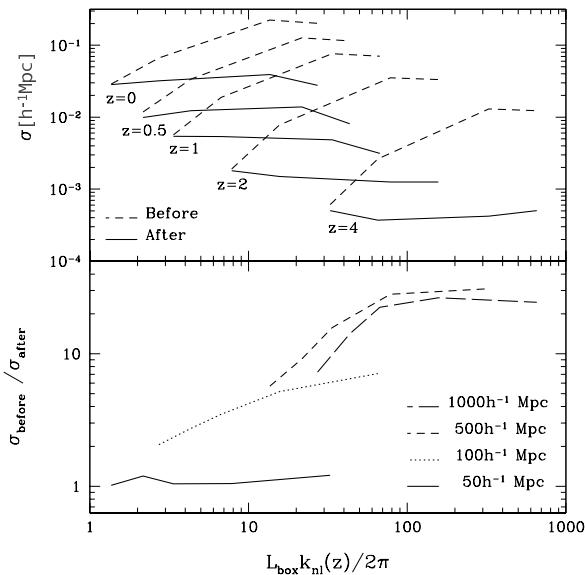


Figure 1. Top panel: 1D rms difference in the rescaled particle positions for the full (solid) and reduced (dashed) AW10 method. The lines show the variation in the errors for fixed redshifts, as a function of the simulation box size; the pairs of solid and dashed lines correspond to different redshifts. Bottom panel: ratio between the 1D rms difference in the particle positions obtained after applying the reduced and full AW10 method; each line shows the results when fixing the box size but allowing the redshift to vary (box sizes are shown in the figure key). The x-axis shows the product of the box side of simulations L_{box} and the non-linear limit mode $k_{nl}(z)$.

simulation was run varying both the Ω_m and σ_8 parameters by $+1\sigma$. The relevant parameters for both sets of simulations are detailed in Table 1. The range of masses over which the minimisation is performed is the same as for the *A* and *B* simulations, in order to mimic the accuracy of this procedure for a simulation with the particle mass resolution of the Millennium Simulation.

All the simulations were evolved from their initial redshifts using the public version of GADGET2 (Springel 2005). The initial conditions were constructed using the public code GRAFIC2 (Bertschinger 2001) and use exactly the same random seed in all cases. The halo catalogues, including substructure identification and merger histories, were constructed using the SUBFIND algorithm explained in detail in Springel et al. (2001). The linking length parameter used is equal to 0.17 times the mean interparticle separation and considering only groups with at least 10 particles. The simulations outputs consist of 100 steps equally spaced in $\log(a)$ between $z = 20$ to $z = 0$.

In the process of obtaining the *B* simulations we search for the scaling factor s and final redshift z_A^f that provide the best fit between the actual linear fluctuation amplitudes via the minimisation of Eq (3). These are shown in Table 2 for each simulation (see the Table caption for the mass ranges used in the minimisation). The table also shows the minimum value of δ_{rms}^2 obtained.

2.2 When is it acceptable to use the reduced AW10 method?

In this subsection we perform a test that will allow us to infer in which cases it is acceptable to ignore the quasi-linear correction of long wavelength contributions. This test consists on implementing the reduced and full algorithm and studying the resulting differences as a function of the boxsize of the simulation, and of the redshift. The latter is done since $k_{nl}(z)$ depends on the redshift.

We use two sets of four 25^3 particle simulations specific for this subsection with reference simulation boxesizes $L_{\text{box}} = 50, 100, 500$ and $1000 h^{-1}\text{ Mpc}$. The first set has the same cosmological parameters than simulation *A* and the second one has the parameter set of simulation *B*. We will analyse the $z = 0, 0.5, 1, 2$ and 4 outputs.

The top panel of Figure 1 shows the 1D rms difference in the particle positions, for the case where the AW10 method is applied in full and in its reduced version (solid and dashed lines, respectively); each line shows the change in the error as the simulation box increases to the right. As can be seen, for simulations of $\sim 50 h^{-1}\text{ Mpc}$ a side, the reduced rescaling algorithm produces positions that are as precise as those obtained using the full method.

The bottom panel shows the ratio between the 1D rms differences obtained from the reduced and full AW10 methods for different values of the simulation box, as the redshift increases to the right. In this case it can be clearly seen that the small simulation box (with side $\sim 2\pi/k_{nl}(z=0)$) shows almost a unit ratio for all the redshifts explored.

The fact that the large-scale correction is not important in simulations of $\sim 50 h^{-1}\text{ Mpc}$ a side is expected since, by construction, the displacement fields are smooth on modes larger than k_{nl} . Note that this correction depends on the cosmologies selected for the simulations, and therefore the results shown in this subsection are only presented as a qualitative example.

3 RESULTS

In this section we apply the rescaling to the simulations presented above, and perform tests on the recovery of the properties of individual haloes, including their detailed growth histories. We also study the recovery of statistical properties of the global population of haloes. From this point on we apply the reduced version of the AW10 algorithm to dark matter haloes in all the simulations except *Bbig*, to which the full method is applied since has a boxsize for which the quasi-linear correction is important.

3.1 Comparison of Individual Halo Properties

We compare the properties of haloes of at least 50 particles in the numerical simulations run with the desired cosmological parameters (*B* simulations, the reference model) to the haloes from simulations *A* (with a different cosmology) rescaled to the cosmology of the *B* simulations (rescaled haloes). These haloes are allowed any number of member particles (at least 10). In order to do this we match haloes that share the largest percentage of particles via their GADGET2 identifying particle ID (referred to as matched haloes);

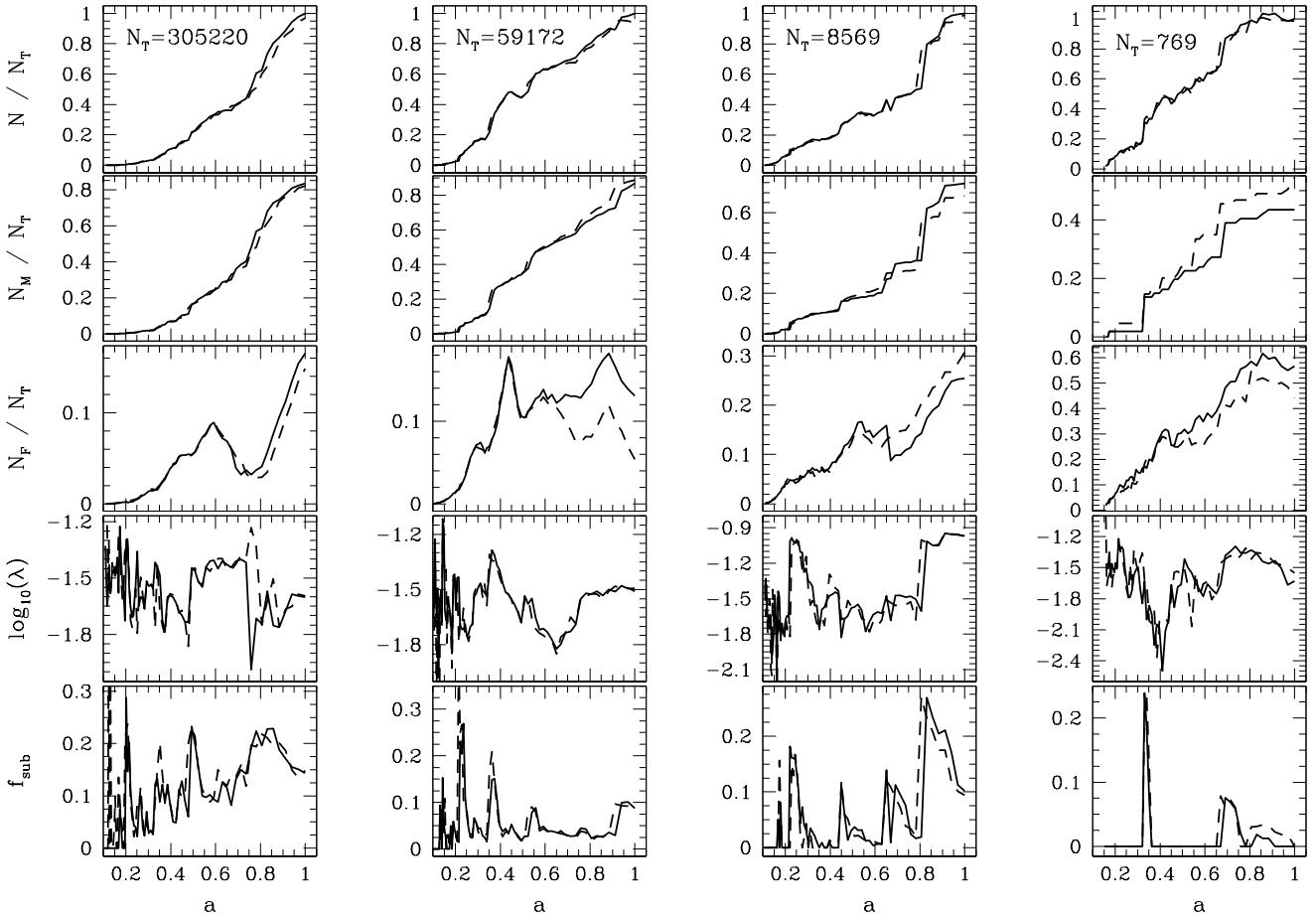


Figure 3. Comparisons between individual haloes belonging to the rescaled and reference catalogues. The number of particles per halo is shown in the upper panels. From top to bottom the panels show the cumulative accretion history, cumulative accretion via mergers, cumulative accretion of individual particles, the evolution of the spin parameter and the fraction of mass in substructures, all as a function of the expansion parameter a . The solid lines show the true evolution; the dashed lines are for the evolution in the rescaled catalogue. The haloes correspond to the 1st, 10th, 100th, and 1000th largest haloes in simulation B and their FoF masses are 4.2×10^{14} , 8.2×10^{13} , 1.2×10^{13} and $1.1 \times 10^{12} h^{-1} M_\odot$ respectively.

this can be done as the initial conditions are constructed using the same random seed. The percentage of haloes in the B simulations that have a matched halo in the A simulation is higher than 99% in all cases.

We compare the properties of matched haloes in Figure 2. The number of particles of the recovered haloes shows a slight underestimation of < 5 percent, and a clear increase in the dispersion that decreases with halo mass and becomes 10 percent for haloes of ~ 200 particles. Results for a lower resolution simulation (*Blow*) in green, and for the simulation with ~ 100 times larger volume (*Bbig*) in red are also shown. As can be seen, neither the lower resolution nor the larger amplitude over which the minimisation is done affect the precision of the recovered number of particles per halo for haloes with > 200 particles. The catalogue of simulation *Bbig* shows a slightly broader range of recovered number of particles for haloes with less particles, but it is not clear whether this is a result of noise or the larger box of the simulation. The results for the *Blow* and *Bbig* simulations are also similar to those of the B simulation in most of the comparisons shown in the other panels of this figure. Therefore, in order to improve clarity, in the remaining panels we will only show the results for the B simulation except for the

cases where there are noticeable differences with resolution or box size.

As can be seen the positions of haloes are well recovered to a precision of $0.1 h^{-1} \text{Mpc}$ for the catalogue of simulation A (difference between percentiles 10 and 90 and the median). The velocities show relative differences of less than ~ 5 percent, and small changes in the direction of the velocity vector with a mode of 1 degree. As can be seen in the figure, the velocities tend to be biased high for $v < 350 \text{ km/s}$, and biased low for $v > 700 \text{ km/s}$, although always below a 5 percent difference. The precision in the recovered halo positions and velocities are consistent with those reported by AW10 in their Figure 8 using the full implementation of the algorithm.

The larger box (blue line) produces an increase in the bias at low peculiar velocities but removes it at the large velocity end. Also, the amplitude of the differences between velocities increases more rapidly for low velocities in the catalogue of the large volume simulation.

We compute the virial mass of haloes as the mass inside a sphere that contains an average virial overdensity of $\delta_{vir}(z, \Omega_m) \approx 18\pi^2(1 + 0.4093x^{2.7152})(1 + x^3)^{-1}$, where $x = (1/\Omega_m(z=0) - 1)^{1/3}(1+z)^{-1}$ (Nakamura & Suto

1997). Since the methodology we apply does not involve re-identifying the haloes using dark-matter particles, the virial mass for the rescaled haloes resulting from applying Equation (6) is underestimated by a ~ 10 percent (panel E, dotted line). However, by considering a correction due to the difference in $\delta_{vir}(z, \Omega_m)$ which is 15 percent higher for the cosmology of the recovered haloes, we are able to reduce the discrepancy in the virial masses to below 5 percent. To do this correction we assume a NFW dark matter density profile (Navarro, Frenk & White 1997) for each halo, compute the integral of the assumed density profile above the corresponding virial overdensity in each cosmology using a concentration parameter given by Bullock et al. (2001), and then multiply the rescaled mass by the ratio between the integrated mass in the target and base models.

We also explored the effect of adopting different concentrations on the recovered virial masses (Figure 2). We used three different concentrations, (i) the proxy v_{max}/v_{vir} (where v_{max} is the maximum circular velocity of the halo and v_{vir} is the circular velocity at the virial radius), (ii) the recipe given by Bullock et al. (2001), and (iii) from the fit provided by Zhao et al. (2008). Our results indicate a negligible effect on the mass, which neither improves nor diminishes the level of agreement between rescaled and reference masses, independently of the definition of concentration applied.

Regarding the recovery of the concentrations of the haloes in the target cosmology, the individual concentrations defined as v_{max}/v_{vir} are affected by a bias of ~ 10 percent with a strong dependence with the concentration in the reference cosmology. On the other hand, the concentrations computed using the mass dependence reported by Bullock et al. (2001) show a negligible difference between rescaled and reference cosmologies, with no dependence on the halo virial mass. The Zhao et al. (2008) concentrations are obtained using their dependence on the cosmological time when the main progenitor acquired 4 percent of the final FoF halo mass. This quantity can be computed using the information in the halo merger trees; however, due to the numerical resolution of our simulation, this quantity can only be calculated for the ~ 2000 most massive haloes. For these haloes, the concentration shows no offset between the rescaled and target simulations but its dispersion is larger than for the Bullock et al. (2001) concentrations, particularly for low c values.

Even though the recovery of individual properties shows some biases and non-negligible scatter, the quantification of these effects can be used to gauge what studies can be performed with rescaled simulations and the statistics that are expected to be affected by these uncertainties.

Our findings on the accuracy of the properties of haloes are consistent with those found by AW10, which indicates that applying the full method to particles or haloes produces only small differences. We confirm this estimate using our simulations, where we only find negligible differences in the resulting accuracy when using individual particles.

3.2 Examples of halo evolution

For a more detailed comparison between the rescaled and reference halo catalogues, we also compare the detailed histories of individual friends-of-friends (FoF) haloes.

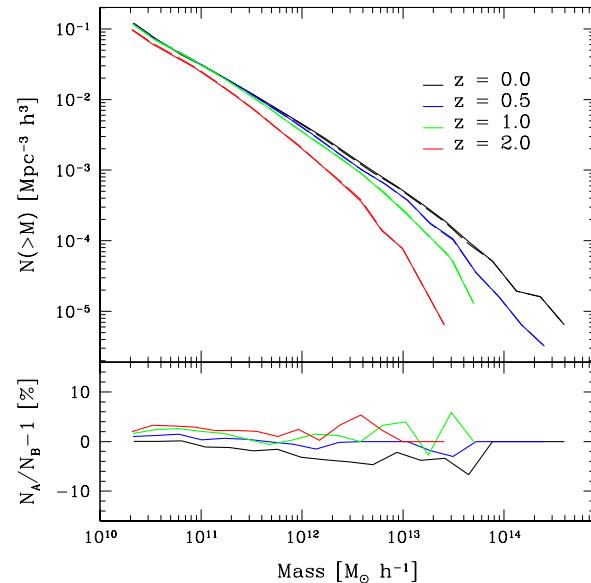


Figure 4. Mass function for rescaled and reference FoF halo catalogues (solid and dashed lines, respectively) for different redshifts (different colours as indicated in the figure key). The lower sub-panel shows the relative difference between the rescaled and reference catalogues, in the same colour scheme.

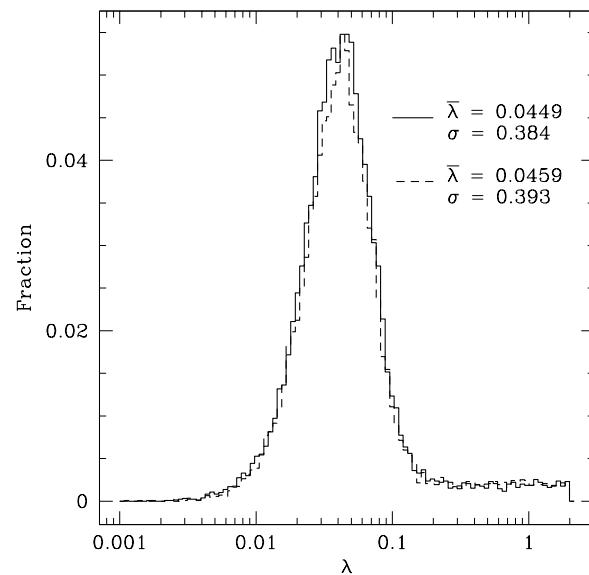


Figure 5. Distribution of halo spin parameters for the rescaled (dashed lines) and reference (solid) catalogues. The mean value and standard deviation of the log-normal distributions are shown in the legend.

We choose the first, 10th, 100th and 1000th most massive haloes and show them in Figure 3. The panels show from top to bottom, the accretion history, the accretion via mergers and via smooth, individual particle infall, the evolution of the dimensionless spin parameter, and the fraction of mass in substructures. The number of particles per halo decreases to the right, and range from FoF masses of

$4.2 \times 10^{14} h^{-1} M_\odot$ to $1.1 \times 10^{12} h^{-1} M_\odot$. Regardless of the halo mass, the individual accretion histories, spin parameter evolution, and fraction of mass in substructures, are remarkably well recovered. The only noticeable differences are shown in the mass accretion via mergers and individual infall of particles, which cancel out as when one is overestimated the other compensates, showing the effects from confusion between accretion of small mass haloes and infalling field particles. This effect can be appreciated more clearly in the mass accretion history of the 10th largest halo in Figure 3.

These results indicate that individual merger trees should be reasonably suitable for rescaling and later use for semi-analytic type galaxy formation modeling. The history of mass accretion via mergers would ensure a 10 percent accuracy in the population of satellite galaxies, whose added stellar masses should be even more precise (according to the well recovered fraction of mass in substructures), and the accurate evolution of the spin parameter should ensure reasonable estimates of galaxy disc sizes.

3.3 Statistical properties of haloes and their growth histories

If the rescaled simulation is used to produce statistics of large populations, either of dark-matter haloes or simulated galaxies obtained via semi-analytic techniques, it is necessary to estimate how they are affected by the rescaling process.

We first study the mass function of dark-matter haloes. Figure 4 compares the mass functions of the rescaled and reference haloes (dashed and solid lines, respectively) for different redshifts from $z = 0$ to $z = 2$. The agreement is excellent with only minor differences. The relative differences shown in the lower panel indicate that there is a mild, although clear trend that starts as an underestimation of abundances at $z = 0$ but then tends to overestimate them by larger amounts at higher redshifts. At low redshift the rescaled catalogue shows a slight underestimation of the number density of haloes which is very mild at low masses $M \sim 10^{10} h^{-1} M_\odot$ at the 1 percent level, but increases to a 5 percent lower abundances at $M \sim 10^{13.5} h^{-1} M_\odot$. The highest redshift shown in the figure indicates a flat overestimation of abundances of a $\sim 3 - 5$ percent.

These effects are small in comparison to the current precision of measurements of the mass function in clusters of galaxies (Gladders et al. 2007) and galaxy luminosity functions (for example in the SDSS, see Blanton et al. 2003) which, in simulations, are highly influenced by the underlying mass function of dark-matter haloes (see for instance Cole et al. 2000). However, the determination of a power spectrum from a mass function requires high precision measurements (Sánchez, Padilla & Lambas 2002) and the accuracy of the rescaling technique could be important in this case.

The spin parameter influences the resulting properties of disc galaxies in semi-analytic models (Cole et al. 2000; Lagos, Padilla & Cora 2009; Tecce et al. 2010). Therefore it is important to check whether the rescaling of a numerical simulation could produce biased distributions of spin parameters. Figure 5 shows these distributions for the rescaled and reference catalogues (dashed and solid lines, respectively), and as can be seen, the distributions are consistent with one

another. The distributions of disc sizes would therefore be expected to be reliable in the rescaled simulation. This is also in agreement with the evolution of the spins of individual matched rescaled and reference haloes shown in Figure 3, which are consistent with one another at all times.

Given that the stellar population of a galaxy is highly dependent on the mass accretion history of a halo, either via mergers or smooth infall, a statistical study of the merger trees of haloes can also be used to test the adequacy of the rescaling. Figure 6 shows the differences between rescaled and reference models for four particular characteristics of merger trees. The quantity shown on the top-left panel corresponds to the time when half the total mass of a halo was formed in the tree t^{sa} (i.e. including all the mass in satellites that will later merge with the central halo); by subtracting any cooling and star formation timescales involved, this would correspond to a stellar age for the final galaxy. As can be seen, the differences show a clear peak at 0Gyr and a median at ~ 0.19 Gyr, with a mean width of 0.84Gyr. On a $z = 0$ galaxy population the former difference would produce little effects on the resulting galaxies, but the latter may spuriously broaden distributions of galaxy properties such as colours.

The second statistic, the time since the halo in the main branch of the merger tree attained half of its final mass, t^a (top-right panel of Figure 6), would be related to the time of assembly of the mass of the final galaxy and shows an offset of ~ 0.19 Gyr and a mean width of 0.86Gyr. Similar results are obtained for the time since the last major merger t^{mm} (bottom-left) which should correlate with the time since the last starburst (the actual starburst should be more recent than this quantity when taking into account the dynamical friction that affect galaxies), which shows a similar offset of ~ 0.12 Gyr on average, and a similar distribution width of 0.69Gyr.

Since the timescales of last starburst, star formation, and stellar mass assembly are affected in a similar way, we do not expect important changes on the spectral shape of a galaxy obtained using rescaled merger-trees, with the exception of the effect from a general shift in age towards older populations by ~ 0.2 Gyr. The lower-right panel of Figure 6 shows the histogram of the difference between the number of major mergers undergone by the rescaled and reference haloes between $z = 4$ and 0. As can be seen, the distribution is symmetric with a clear peak at zero (which contains more than 50 percent of the sample), and a maximum difference of 2 or more major mergers for about 10 percent of the haloes.

AW10 presented comparisons between semi-analytic galaxies obtained from rescaled and direct simulations, showing an effect that may be the product of these offsets, the slightly fainter K-band magnitudes (by 0.1 mag) in their rescaled simulations.

The study of the effect of the rescaling of merger trees on the inferred positions of haloes performed in Section 3 showed that these are not badly affected by the process. However, the correlation function could show changes due to the dispersion in the positions shown in the top-left panel of Figure 2. Figure 7 shows the resulting cross-correlation functions from the rescaled and reference haloes. For this test we choose cross-correlations over auto-correlation functions in order to increase the signal-to-noise ratio of our measurements (see for instance Bornancini et al.

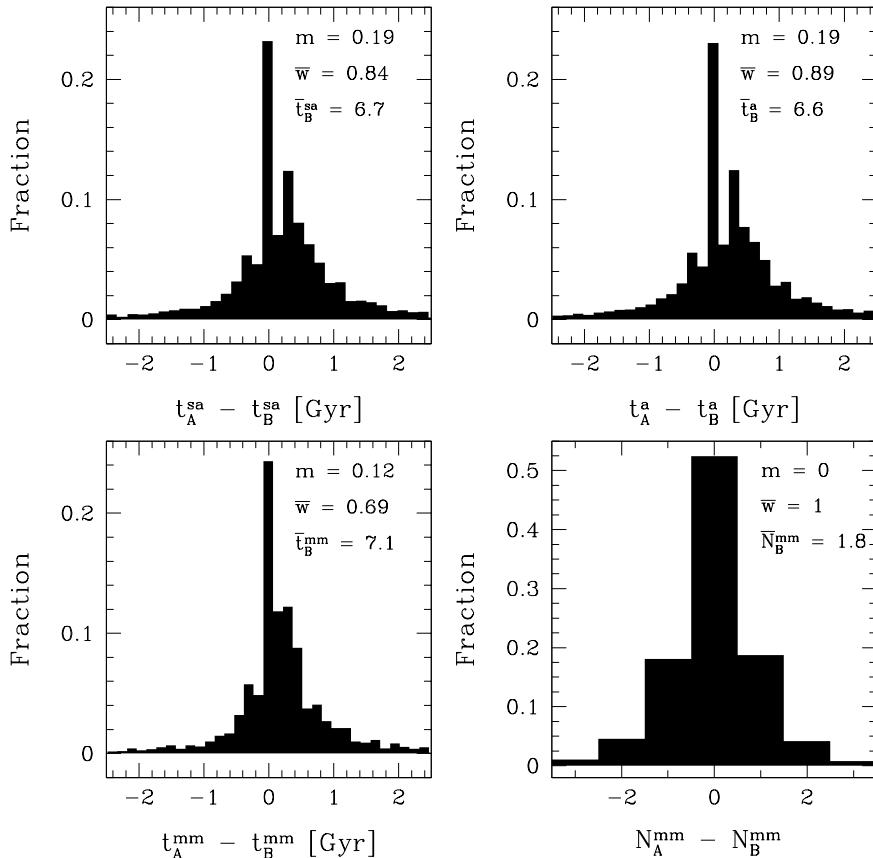


Figure 6. Distributions of differences between merger-tree properties for dark-matter haloes in the rescaled and reference model. Top-left: distribution of the error in the time since the mass added over all the branches of the merger tree reached half the final mass. Top-right: same as top-left for the time since the halo in the main branch of the merger tree attained half its final mass. Bottom-left: same as previous panels for the time since the last major merger (with at least a 0.3 mass ratio). Bottom-right: differences between the number of major mergers for haloes which have undergone at least one major merger. The median m and mean width \bar{w} of the distributions are given in the legend of each panel; we also show the average quantities in the reference cosmology. For the upper panels and the lower-left panel, the units are Gyr.

2006; Lacerna & Padilla 2011). Halos of different lower limits on FoF mass are used as centres, whereas all haloes in the simulations are used as tracers for these measurements. We choose three different lower limits on the mass of centre haloes ($M_C > 10^{12} h^{-1} M_\odot$, $M_C > 10^{13} h^{-1} M_\odot$ and $M_C > 3 \times 10^{13} h^{-1} M_\odot$) such that they bracket the non-linear mass for the cosmology of simulation B , $M_{nl} \sim 10^{13} h^{-1} M_\odot$, around which the bias factor shows a clear increase (for haloes with mass M_{nl} , the bias factor is $b = 1$, for more details see for instance Sheth, Mo & Tormen 2001).

As can be seen, at scales $r > 0.5$, 2.5, and $4 h^{-1} \text{Mpc}$ for masses $M_C > 1 \times 10^{12}$, 1×10^{13} and $3 \times 10^{13} h^{-1} M_\odot$, respectively, the precision of the correlation function of the rescaled haloes is better than 5 percent (see the ratios on the lower sub-panel). Given the offset in the number of particles between the reference and rescaled haloes, we also tested whether using the same equivalent lower limit in number of particles improved this comparison, but find very similar results. This is also the case for the effect of the dispersion between rescaled and reference halo masses (number of particles), since the resulting change is of only a few percent,

which influences the clustering amplitude by factors below the offsets originating from the rescaling procedure.

This level of precision is of the order of that obtained for large surveys at $z = 0$ such as the SDSS (see for instance Zehavi et al. 2004). Furthermore, other effects such as the assembly bias (Gao, Springel & White 2005) are expected to produce variations on the clustering of haloes and, consequently, galaxies to a 10 percent, which indicates that rescaled haloes could be used to test the detectability of this particular effect. However, Wu, Rozo & Wechsler (2008) point out that the precision of future surveys such as the Dark Energy Survey (DES, Tucker et al. 2007) and the Large Synoptic Survey Telescope (LSST, Abell et al. 2009) will be much better than that in the SDSS, and could require new runs with the desired cosmology rather than rescaled halo catalogues.

As a final test, we study the changes introduced by the rescaling of haloes in the fraction of FoF particles in substructures (i.e., not restricting substructures to the sphere contained in one virial radius). Figure 8 shows this fraction as a function of the number of particles per halo, for

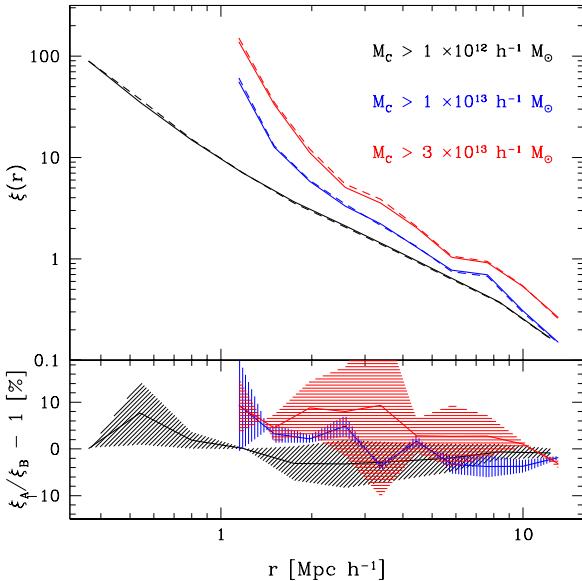


Figure 7. Cross correlation functions between haloes in the rescaled A (dashed lines) and reference B (solid lines) catalogues with different lower limits on FoF mass (different colours shown in the key) and the full population of haloes (with at least 10 particles). The lower sub-panel shows the relative differences between the correlation functions obtained from the rescaled and reference catalogues; the shaded regions show the uncertainties computed using the bootstrap resampling technique.

the rescaled and reference halo catalogues (lines and dots in blue and red, respectively). Both the median relation, and the envelope of 80 percent of the haloes show an excellent agreement between the two catalogues, with differences that only amount to a 10 percent. This statistics is related to different properties of the galaxies inhabiting these haloes. On the one hand, it would influence the variation of the amplitude of clustering as a function of subhalo mass (or galaxy luminosity), since the average host halo mass of a selection of subhaloes can be affected if this fraction is not accurately recovered. On the other hand, this result is in agreement with our earlier claim that we would not expect to find correlated changes in the spectra of semi-analytic galaxies obtained from the rescaled trees. This is due to the fact that the changes in the characteristic timescales of the merger trees are slightly biased but these biases go all in the same direction, otherwise we would expect to find shifts in the fraction of mass in subhaloes. This is also the cause for the fraction of mass in substructures to remain consistent; this can happen if the overall histories are consistent, even if slightly shifted.

4 TESTS ON VARYING THE EXTRAPOLATION BASELINE

The method of adjusting the cosmological parameters of a set of haloes and their assembly histories can be very useful for exploring a cosmological parameter space using halo properties from numerical simulations. In this section we investigate how the agreement between the rescaled and ref-

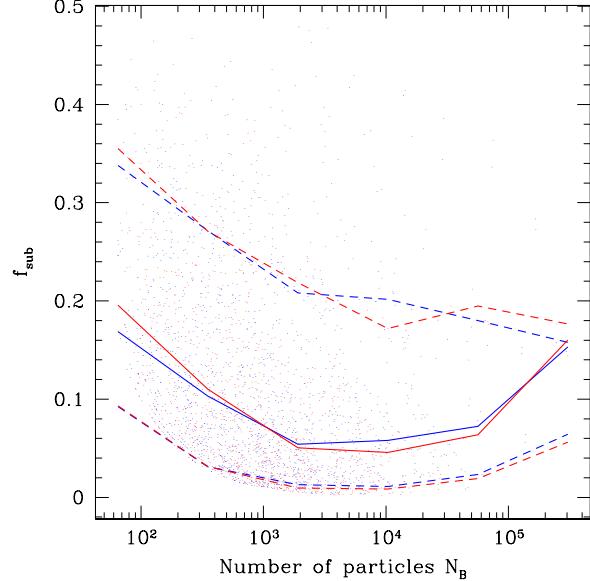


Figure 8. Dependence of the median (solid lines), 10 and 90 percentiles (dashed lines) of the difference between the recovered and true fraction of mass in substructures as a function of the number of particles per halo. Blue and red lines correspond to the rescaled and reference halo catalogues, respectively. The uprising of f_{sub} at small masses may correspond to a resolution effect.

erence haloes degrades as the baseline of the extrapolation on the parameter space varies.

We choose to vary only two cosmological parameters for this test, the matter density parameter Ω_m , and the amplitude of fluctuations σ_8 . In this case we take as a starting point the parameters of simulation *Blow*,¹ which are in agreement with the latest constraints on the cosmology (Jarosik et al. 2010), and explore the resulting rescaled haloes for the Bo_{-2} , Bo_{-1} , Bo_{+1} , Bo_{+2} simulations, which maintain the same cosmological parameters as simulation *Blow* but vary Ω_m between $\Omega_m - 2\sigma$, and $\Omega_m + 2\sigma$ with $\sigma = 0.03$, the simulations Bs_{-2} , Bs_{-1} , Bs_{+1} , Bs_{+2} , which with respect to *Blow* only change σ_8 between the values $\sigma_8 - 2\sigma$, and $\sigma_8 + 2\sigma$ with $\sigma = 0.03$, and a simulation in which both, Ω_m and σ_8 are increased in one standard deviation, $Bo_{+1}s_{+1}$ (details on the simulation parameters are shown in Tables 1 and 2). In all cases, the value of σ is similar to the standard deviation in these parameters from the WMAP7 results. Since we are using simulations with small boxes for this test, we only apply the reduced version of the AW10 algorithm in this section.

Figure 9 shows average and median quantities that summarise the accuracy of the rescaled haloes. We include the variation in the number of particles per halo averaged over all haloes with 200 to 1000 particles and the 10 and 90 percentiles of the distribution (top-left), the average relative variation in the cumulative mass function for masses

¹ The choice of small box and low resolution is justified by the precision of the rescaled haloes in the low resolution simulation *Blow* shown in Figure 2, and owes to the large number of simulations involved in this analysis.

in the range $11 < \log_{10}(M/h^{-1}M_\odot) < 13$ and the dispersion of the distribution (top-right), the average variation in the cross-correlation function between haloes of FoF masses $> 10^{12}h^{-1}M_\odot$ against the full halo population, in the range of scales $1 < r/h^{-1}\text{Mpc} < 10$ and its dispersion (lower-left), and the median difference in the assembly ages of rescaled and reference haloes and the 10 and 90 percentiles of the distribution (lower-right). In all cases, the x -axis shows the number of WMAP7 standard deviations from the *Blow* parameters. As can be seen, the effect from changing the matter density parameter is in all cases more important than varying the amplitude of fluctuations σ_8 . We also show the resulting biases and uncertainties in the solid and open green symbols when changing both Ω_m and σ_8 by $\pm 1\sigma$, which follows the expected tendency and results in larger biases and dispersions than varying only one parameter at a time.

These statistics are chosen since they are representative of important variations in the galaxy population that can be obtained via semi-analytic models (for instance via the abundance of haloes or the assembly time-scales), or variations in constraints on cosmological models via clustering and abundance measurements.

The biases in the number of particles per halo and the number of haloes shown in the top panels are correlated, and consist on underestimates (overestimates) for lower (higher) values of either Ω_m and σ_8 . This is due to the fact that a bias in the number of particles, which does not depend on the halo mass as was seen in the comparison between simulations *A* and *B*, *Alow* and *Blow* and *Abig* and *Bbig* (cf. Figure 2, although an underestimation of the number of particles per halo is present), is reflected in the mass function. The change from simulation *A* to *B* is approximately equivalent to -3σ and $+1\sigma$ in σ_8 and Ω_m , respectively, which would correspond to $\sqrt{1^2 + 3^2}\sigma \sim -3\sigma$ in the top-left panel of Figure 9, where the number of particles per halo in rescaled catalogues is underestimated, showing the consistency of our analysis.

The lower-left panel of Figure 9 shows that the amplitude of the correlation function shifts from an overestimation when the parameters are lowered, to an underestimation when going to higher values of either Ω_m or σ_8 (with a larger variation in the former). Finally, the assembly times shown on the lower-right show almost no biases, but a slight tendency towards underestimating this quantity for positive σ values on either σ_8 or Ω_m (and stronger for the latter).

The increase in number of particles per halo, the space density of haloes, and the decrease in the amplitude of the correlation function (and possibly in the assembly time-scales) responds to the decreasing values of the rescaling parameter s . This can be seen in Figure 10 where we show in a colour gradient that, for a fixed value of Ω_m (for instance, that of the *Blow* simulation, shown by the vertical line), a higher final value of matter density parameter requires $s < 1$. This figure also serves the purpose of allowing to rescale the expected biases shown in Figure 9 to a different base cosmology. The white diagonal lines are contours of the equal s values that satisfy $\Delta\Omega_m = 0.03$ at the *Blow* cosmological parameter set, and show that for lower values of base cosmology Ω_m , the biases of Figure 9 would only be obtained for smaller values of $\Delta\Omega_m$. As this result is independent of the base and desired values of σ_8 , this shows that a grid of cosmological simulations would need to more densely cover low Ω_m values, roughly following $\Delta\Omega_m = 0.03 \times (\Omega_m/0.27)$,

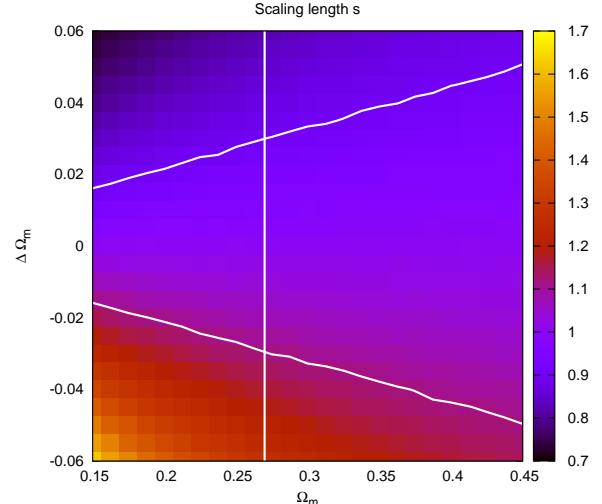


Figure 10. Rescaling parameter s (color gradient) as a function of the starting cosmology parameter Ω_m (x-axis) and $\Delta\Omega_m$ needed to reach the desired cosmology (y-axis). The remaining cosmological parameters are those of the *B* simulation, with the exception of σ_8 which is set to 0.82 for the starting cosmology and 0.73 for the desired one. The white vertical line indicates the position of $\Omega_m = 0.27$, and the diagonal lines represent the contours crossing $\Omega_m = 0.27$ when $\Delta\Omega_m = \pm 0.03$. The scaling length parameter s is not sensitive to changes in σ_8 .

to ensure a stable accuracy in rescalings done along the Ω_m parameter, for any value of σ_8 . As for changes in σ_8 , as this will entail changing the output redshift adopted for the rescaling rather than introducing a new dimension in the grid of simulations, this will not affect the design of the grid.

These results can be used to set the maximum difference on either Ω_m or σ_8 that will be allowed when adopting the AW10 method to rescale dark-matter haloes from a given cosmology into different ones. This result is also expected to hold for larger simulation boxes as long as the full rescaling algorithm is applied.

5 CONCLUSIONS

We have tested the use of the full and a reduced form of the Angulo & White (2010) method to change the cosmology of a catalogue of simulated haloes with the aim to test its application to a cosmological parameter search using semi-analytic galaxies. Our tests on the effect of applying it to the haloes in a numerical simulation instead of to the individual particles, shows a dramatic reduction of computational time. The main reason for this particular test is that future planned numerical cosmological simulations will be so large that the storage of their individual particles will be impractical, in which case only halo properties will be available. Even in the event that the approach presented here is applied to a simulation with available individual particles, the CPU time required to rescale DM haloes is orders of magnitude (at least one) smaller than it is required to rescale the particles. If one adds the time consumed in identifying DM haloes and constructing merger trees, the speed up of

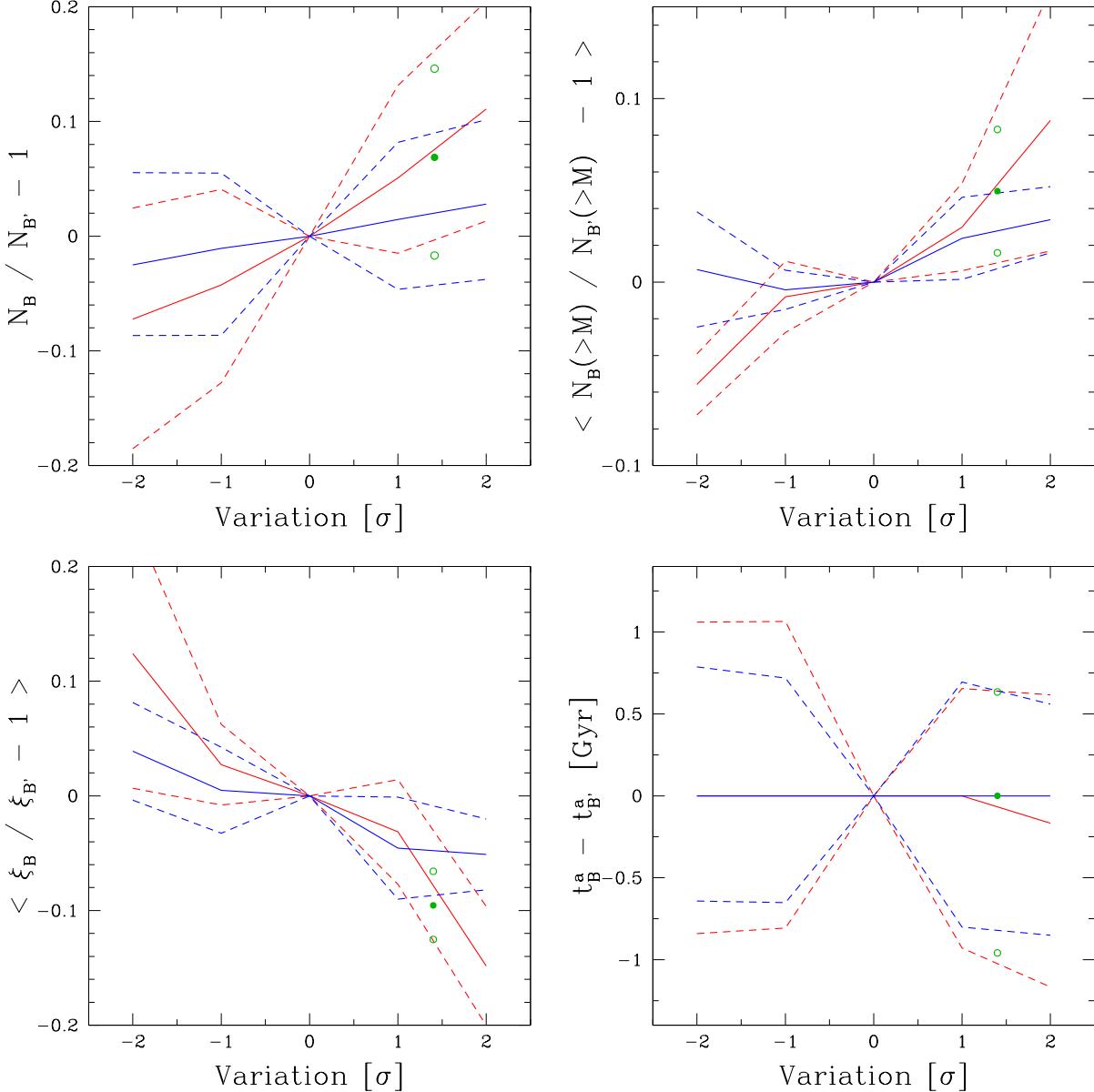


Figure 9. Recovery of halo properties as the baseline for the extrapolation in cosmological parameters extends away from the base values. In all the panels the blue lines correspond to varying the σ_8 parameter, and the red lines to variations in Ω_m . Solid lines show the mean (top-right and bottom-left) or median values (top-left and bottom-right). The dashed lines show the 10 and 90 percentiles of the resulting distributions in the upper-left and lower-right, and the dispersion in the other two panels. The green solid symbol shows the recovery of halo properties for the model in which both, Ω_m and σ_8 are increased by 1σ , shown at $\sqrt{1^2 + 1^2}\sigma$ on the x-axis; the open symbols show the 10 and 90 percentiles, or the dispersion. Top-left: variation in the cumulative number of haloes with masses $11 < \log_{10}(M/h^{-1}M_\odot) < 13$. Top-right: variation in the number of particles per dark-matter halo, for haloes with 200 to 1000 particles. Bottom-left: variation in the correlation functions, averaged on $1 < r/h^{-1}\text{Mpc} < 10$. Bottom-right: difference between the assembly time-scales t^a of rescaled and reference haloes.

the process can reach an improvement of two or three orders of magnitude (R. Angulo, private communication). The reduced form of the rescaling algorithm further reduces the computational time, but at the expense of being limited to small simulation boxes. For the cosmological parameters of simulations A and B (see Table 2), the simulation box side should be $\sim 50h^{-1}\text{Mpc}$ or smaller. In such boxes, the correction for the quasi-linear modes does not further improve the accuracy of the rescaling.

We measured the achieved precision of the rescaled haloes extracted from a given set of cosmological parameters (such as those adopted in the Millennium simulation) by comparing them to those extracted from numerical simulations with the desired final parameters (in a first instance corresponding to a Bolshoi simulation cosmology), constructed using the same initial conditions as the simulation from which the rescaled haloes are taken. We compared individual properties, such as the number of particles per

halo, their peculiar velocities (and variation in the direction of movement), concentration parameters, the mass of the haloes, and their positions. In all cases the precision of the recovered properties is comparable to what is obtained from the full modification of the individual particles in the simulation (AW10). The level of precision is better than 100kpc for the halo positions, and of 5 percent of their peculiar velocities. Both the number of particles per halo, and the virial halo mass are slightly underestimated, by 5 percent.

It should be noticed, though, that in order to obtain this accuracy for the virial mass, it is necessary to perform an additional correction which takes into account the different values of $\delta_{vir}(z, \Omega_m)$ in the rescaled and target cosmologies. The halo concentrations derived using the Bullock et al. (2001) recipe shows the lowest biases between rescaled and reference cosmologies, with no dependence on halo mass.

We tested the effects from lowering the resolution of the simulation by a factor of 8 in the number of particles (for the same total periodic volume), and from increasing by a factor of ~ 100 the volume of the simulation (for the same number of particles). We expect changes in the two cases due to different reasons. In the case of changing the resolution, all the other parameters of the rescaling of the haloes were held fixed. In the change of the simulated volume, the range of masses over which the rescaling function was minimised was shift one order of magnitude to higher masses (see Table 2). Lowering the resolution or increasing the volume do not change the precision in the recovered number of particles as a function of the number of particles in the halo. This implies that a lower resolution results in a lower precision at a fixed halo mass. The lower resolution does not produce any other significant biases. The increase of the volume does not result in a higher uncertainty in the recovered positions since the quasi-linear correction of long wavelength contributions in the AW10 method properly takes into account the different quasi-linear theory modes in the two models.

Furthermore, we also explored the differences arising in the detailed history of growth of haloes, and find only mild displacements in particular events such as major mergers. In order to do this we checked haloes corresponding to the 1st., 10th, 100th and 1000th most massive objects in the simulations, and the resulting agreement is independent of the halo mass, at least down to masses $M \simeq 8 \times 10^{11} h^{-1} M_\odot$. These results indicate that rescaled haloes and halo histories can even be used independently with little effects on the resulting galaxy population.

A direct application of this method consists on changing the cosmological background of semi-analytic galaxies. In order to assess the possible systematic biases that such a population of galaxies would present if it was based on a catalogue of rescaled haloes, we studied the offset in several measures of the halo merger histories, on the halo mass function and their clustering properties, on their spin parameter distributions, and in the number of major mergers experienced by the haloes. Neither the mass function nor the correlation amplitudes show important variations (< 5 percent) within the range of masses explored, which is also the case for the distribution of spin parameters or frequency of major mergers. The only possible important discrepancy comes from a systematic offset in the ages of merger trees (assembly, star formation and last starburst), which are all biased toward smaller ages. Given that these are correlated,

the relative shapes of the spectral energy distributions of galaxies of different types will remain almost unchanged, but will result in a global change of the population toward slightly bluer colours.

This method can effectively allow to sample the cosmological parameter space using fully non-linear simulations. Even though semi-analytic models were designed to understand the processes driving galaxy formation and evolution, the possible dependences of galaxy properties on the cosmological parameters could also be used to impose constraints on the latter. To do this one possible approach is to use Monte-Carlo Markov Chain or similar analyses of the cosmological parameter space (see for instance Harker, Cole & Jenkins 2007 and Bower et al. 2010). This would produce chains of parameter sets on which, in principle, a new numerical simulation would need to be run, in which the detection of haloes and merger trees, and a semi-analytic model, would need to be performed and applied. This process becomes much more efficient when using rescaled haloes, which avoid several time-consuming processes. The lowest computational cost is achieved by rescaling haloes in boxes small enough so as to allow the use of the reduced method.

At this point it is important to note that even if there are small systematic effects on the scaled catalogues and merger trees, these are negligible in comparison to the uncertainties in the semi-analytic modeling. Furthermore, given the flexibility of semi-analytic models, it is in principle possible to diminish the difference between rescaled and direct simulations, but in a way that is not related to the scaling formalism; therefore, this particular ability does not help to improve the use of rescalings in a fast search for baryonic and cosmological parameters together.

We showed that in terms of the uncertainties in the matter density parameter Ω_m and amplitude of fluctuations parametrised by σ_8 , the biases on different halo properties increase considerably more in the case of adopting more different values of Ω_m than of σ_8 . These constraints can be used to set limits on how far the extrapolation can be extended in terms of these parameters; this way only a limited number of numerical simulations located in strategically selected grid points would need to be run in order to cover in a continuous way a wide cosmological parameter space.

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REFERENCES

Abell P., et al. (The LSST Collaboration), 2009, The LSST Science Book, astro-ph/arXiv:0912.0201

Angulo R. E., White S. D. M., 2010, MNRAS, 405, 143

Baugh C. M., Lacey C. G., Frenk C. S., Granato G. L., Silva L., Bresca A., Benson A. J., Cole S., 2005, MNRAS, 356, 1191

Bertschinger E., 2001, ApJS, 137, 1

Blanton M. R., Hogg D. W., Bahcall N. A., Brinkmann J., Britton M., Connolly A. J., Csabai I., Fukugita M., Loveday J., Meiksin A., Munn J. A., Nichol R. C., Okamura S., Quinn T., Schneider D. P., Shimasaku K., Strauss M. A., Tegmark M., Vogeley M. S., Weinberg D. H., 2003, ApJ, 592, 819

Bornancini C. G., Padilla N. D., Lambas, D. G., De Breuck C., 2006, MNRAS, 368, 619

Bower R. G., Vernon I., Goldstein M., Benson A. J., Lacey C., Baugh C. M., Cole S., Frenk C. S., 2010, MNRAS, 407, 2017

Bower R. G., Benson A. J., Malbon R., Helly J. C., Frenk C. S., Baugh C. M., Cole S., Lacey C. G, 2006, MNRAS, 370, 645

Bullock J. S., Kolatt T. S., Sigad Y., Somerville R. S., Kravtsov A. V., Klypin A. A., Primack J. R., Dekel A., 2001, MNRAS, 321, 559

Cole S., Aragon-Salamanca A., Frenk C. S., Navarro J. F., Zepf S. E., 1994, MNRAS, 271, 781

Cole S., Lacey C. G., Baugh C. M., Frenk C. S., 2000, MNRAS, 319, 168

Croton D. J., Springel V., White S. D. M., De Lucia G., Frenk C. S., Gao L., Jenkins A., Kauffmann G., Navarro J. F., Yoshida N., 2006, MNRAS, 365, 11

Dunkley J., Komatsu E., Nolta M. R., Spergel D. N., Larson D., Hinshaw G., Page L., Bennett C. L., Gold B., Jarosik N., Weiland J. L., Halpern M., Hill R. S., Kogut A., Limon M., Meyer S. S., Tucker G. S., Wollack E., Wright E. L., 2009, ApJS, 180, 306.

Gao L., Springel V., White S. D. M., 2005, MNRAS, 363, 66

Gladders M.D., Yee H.K.C., Majumdar S., Barrientos L.F., Hoekstra H., Hall P.B., Infante L., 2007, ApJ, 655, 128

Harker G., Cole S., Jenkins A., 2007, MNRAS, 382, 1503

Hinshaw G., Weiland J. L., Hill R. S., Odegard N., Larson D., Bennett C. L., Dunkley J., Gold B., Greason M. R., Jarosik N., Komatsu E., Nolta M. R., Page L., Spergel D. N., Wollack E., Halpern M., Kogut A., Limon M., Meyer S. S., Tucker G. S., Wright E. L., 2009, ApJS, 180, 225

Jarosik N., Bennett C. L., Dunkley J., Gold B., Greason M. R., Halpern M., Hill R. S., Hinshaw G., Kogut A., Komatsu E., Larson D., Limon M., Meyer S. S., Nolta M. R., Odegard N., Page L., Smith K. M., Spergel D. N., Tucker G. S., Weiland J. L., Wollack E., Wright E. L., 2011, ApJS, 192, 14

Klypin A., Trujillo-Gomez S., Primack J., 2010, arXiv:astro-ph/1002.3660v4

Komatsu E., Dunkley J., Nolta M. R., Bennett C. L., Gold B., Hinshaw G., Jarosik N., Larson D., Limon M., Page L., Spergel D. N., Halpern M., Hill R. S., Kogut A., Meyer S. S., Tucker G. S., Weiland J. L., Wollack E., Wright E. L., 2009, ApJS, 180, 330

Lacerna I., Padilla N. D., 2011, MNRAS, 412, 1283

Lagos C. D. P., Padilla N. D., Cora S. A., 2009, MNRAS, 395, 625

Nakamura T., T., Suto Y., 1997, Progress of Theoretical Physics, 97, 49

Navarro J. F., Frenk C. S., White S. D. M., 1997, ApJ, 490, 493

Percival W. J., Nichol R. C., Eisenstein D. J., Frieman J. A., Fukugita M., Loveday J., Pope A. C., Schneider D. P., Szalay A. S., Tegmark M., Vogeley M. S., Weinberg D. H., Zehavi I., Bahcall N. A., Brinkmann J., Connolly A. J. Meiksin A., 2007, ApJ, 657, 645

Rozo E., Rykoff E. S., Evrard A., Becker M., McKay T., Wechsler R. H., Koester B. P., Hao J., Hansen S., Sheldon E., Johnston D., Annis J., Frieman J., 2009, ApJ, 699, 768.

Sheth R. K., Mo H. J., Tormen G., 2001, MNRAS, 323, 1

Sánchez A., Padilla N., Lambas D.G., 2002, MNRAS, 337, 161

Sánchez A. G., Baugh C. M., Percival W. J., Peacock J. A., Padilla N. D., Cole S., Frenk C. S., Norberg P., 2006, MNRAS, 366, 189

Sheth R. K., Mo H. J., Tormen, G., 2001, MNRAS, 323, 1

Spergel D. N., Verde L., Peiris H. V., Komatsu E., Nolta M. R., Bennett C. L., Halpern M., Hinshaw G., Jarosik N., Kogut A., Limon M., Meyer S. S., Page L., Tucker G. S., Weiland J. L., Wollack E., Wright E. L., 2003, ApJS, 148, 175.

Springel V., White S. D. M., Tormen G., Kauffmann G., 2001, MNRAS, 328, 726.

Springel V., 2005, MNRAS, 364, 1105

Springel V., White S. D. M., Jenkins A., Frenk C. S., Yoshida N., Gao L., Navarro J., Thacker R., Croton D., Helly J., Peacock J. A., Cole S., Thomas P., Couchman H., Evrard A., Colberg J., Pearce F., 2005, Nature, 435, 629

Tecce T., Cora S., Tissera P., Abadi M., Lagos C., 2010, MNRAS, 408, 2008

Tucker D. L., Annis J. T., Lin H., Kent S., Stoughton C., Peoples J., Allam S. S., Mohr J. J., Barkhouse W. A., Ngeow C., Alam T., Beldica C., Cai D., Daues G., Plante R., Miller C., Smith C., Suntzeff N. B., 2007, ASPC, 364, 187.

Wu H., Rozo E., Wechsler R. H., 2008, ApJ, 688, 729

Zehavi I., Weinberg D. H., Zheng Z., Berlind A. A., Frieman J. A., Scoccimarro R., Sheth R. K., Blanton M. R., Tegmark M., Mo H. J., Bahcall N. A., Brinkmann J., Burles S., Csabai I., Fukugita M., Gunn J. E., Lamb D. Q., Loveday J., Lupton R. H., Meiksin A., Munn J. A., Nichol R. C., Schlegel D., Schneider D. P., SubbaRao M., Szalay A. S., Uomoto A., York D. G., 2004, ApJ, 608, 16

Zhao D. H., Jing Y. P., Mo H. J., Börner G., 2008, ApJ, 707, 354

Zheng Z., Tinker J. L., Weinberg D. H., Berlind A. A., 2002, ApJ, 575, 617